GAS in Protoplanetary Systems (GASPS)

Article - January 2010

52 authors, including:

Jonathan P. Williams  
University of Hawai'i at Mānoa  
244 PUBLICATIONS  6,438 CITATIONS

Gwendolyn Meeus  
Universidad Autónoma de Madrid  
137 PUBLICATIONS  2,548 CITATIONS

Basmah Riaz  
University of Hertfordshire  
49 PUBLICATIONS  494 CITATIONS

Göran Sandell  
Universities Space Research Association  
217 PUBLICATIONS  3,306 CITATIONS

Some of the authors of this publication are also working on these related projects:

- Fourier-Kelvin Stellar Interferometer [View project]
- Atmospheric Bio-signatures [View project]
Letter to the Editor

GAS in Protoplanetary Systems (GASPS)*

I. First results


(Affiliations are available in the online edition)

Received 31 March 2010 / Accepted 9 May 2010

ABSTRACT

Context. Circumstellar discs are ubiquitous around young stars, but rapidly dissipate their gas and dust on timescales of a few Myr. The Herschel Space Observatory allows for the study of the warm disc atmosphere, using far-infrared spectroscopy to measure gas content and excitation conditions, and far-IR photometry to constrain the dust distribution.

Aims. We aim to detect and characterize the gas content of circumstellar discs in four targets as part of the Herschel science demonstration phase.

Methods. We carried out sensitive medium resolution spectroscopy and high sensitivity photometry at λ ~60–190 μm using the Photodetector Array Camera and Spectrometer instrument on the Herschel Space Observatory.

Results. We detect [OI] 63 μm emission from the young stars HD 169142, TW Hydrae, and RECX 15, but not HD 181327. No other lines, including [CII] 158 and [OI] 145, are significantly detected. All four stars are detected in photometry at 70 and 160 μm. Extensive models are presented in associated papers.

Key words. stars: pre-main sequence – protoplanetary disks – infrared: stars

1. Introduction

Circumstellar discs of gas and dust are the environments in which planets form; thus, studies of the evolution of these discs will inform our understanding of the formation of gas giants. The Herschel Space Observatory (Pilbratt et al. 2010), with its facilities for sensitive spectrometry and photometry, allows for significantly deeper searches for circumstellar gas than were previously possible, providing tighter constraints on the timescale of giant planet formation and the mechanisms of circumstellar disc evolution.

The GAS in Protoplanetary Systems Herschel open time key project (GASPS, P.I.W.R.F. Dent, see Dent et al., in prep.) aims to characterize the evolution of gas and dust in circumstellar discs. Using the Photodetector Array Camera and Spectrometer (PACS, Poglitsch et al. 2010) integral field unit spectrometer and bolometer array, we are carrying out a 400 hour survey of 240 young stars for gas line emission and sensitive far-IR photometry. We focus on three primary emission lines, [OI] 63.184 μm, [OI] 145.525 μm, and [CII] 157.741 μm, as well as several emission lines of H2O and CO and broadband photometry at 70 and 160 μm to measure dust emission. Our survey covers young stars at ages from 1–30 Myr, in the Taurus, Upper Scorpius, η Chamaeleontis, TWA, Beta Pictoris, and Tucana Horologium associations, as well as HAeBe stars at a variety of ages 0.1–30 Myr, and with masses ranging from ~0.5–4 M⊙. We chose systems known from near-IR and millimeter photometry and accretion indicators to have discs in a variety of states, from those still accreting to debris discs, as well as a sample of stars showing no signs of a disc. We focus on optically-visible stars with low extinction and small envelope masses, i.e. SED Class II-III.

At a basic level, observations of [CII] and [OI] will provide a determination of the presence of gas. Comparison to photodissociation models (e.g. Kaufman et al. 1999) allows for the use of [CII] and [OI] emission to place constraints on the density and temperature structure of circumstellar discs. Such observations will help illuminate disparities between disc mass estimates based on sub-mm continuum measurements of dust emission and gas tracers such as CO. However, atomic C and O have emission lines in the far-IR that are largely unobservable from the ground. While ISO allowed for studies of atomic emission from a few bright sources (e.g. Lorenzetti et al. 1999), the Herschel PACS instrument has 2 orders of magnitude better point-source sensitivity than the ISO Long Wavelength Spectrograph, opening up the study of circumstellar disc gas to hundreds more sources.

* Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.
To aid with the interpretation of our PACS line and continuum data, we have interfaced tools for calculating dust and line radiative transfer (MCFOST, Pinte et al. 2006) and gas thermo-chemistry (ProDiMo, Woitke et al. 2009). We have used these codes to produce a grid of 300 000 disc models, the Disc Evolution with Near Theory (DENT, Woitke et al. 2010).

In this paper, we present the first results of science demonstration phase observations of a Herbig Ae star (HD 169142), two T-Tauri stars (RECX 15 in η Cha and TW Hya), and a debris disc system (HD 181327). All have ages in the range 6–12 Myr, and bear discs that have undergone some evolution. Near/mid-IR photometry and mm photometry have indicated disc clearing or grain growth in the inner discs of HD 169142 and TW Hya (e.g. Grady et al. 2007; Hughes et al. 2007). Mass accretion rate estimates range from \(<7 \times 10^{-10}\) for HD169142 (Grady et al. 2007) to \(5 \times 10^{-10}–2 \times 10^{-9}\) \(\text{M}_\odot/\text{yr}\) for TW Hya (Muzerolle et al. 2000; Herczeg et al. 2004); and \(~10^{-9}\) \(\text{M}_\odot/\text{yr}\) for RECX 15 (Lawson et al. 2002).

In Sect. 2, we describe our observations and data reduction, and present our results. We present a preliminary discussion of the implications for the states of these discs in Sect. 3.

2. Observations and results

For our SDP targets, we carried out both spectroscopy and photometry with the PACS instrument on Herschel. For each target, we carried out a 1669 s PacsLineSpec observation, a 5150 s PacsRangeSpec observation, and either a 133 s PacsPhoto scan map (RECX 15 and TW Hya) or 159 s PacsPhoto point source observation (HD 169142 and HD 181327). We carried out an additional 220 s scan map observation of RECX 15, to test observation modes. We carried out the reduction of our observations with the facility reduction tool, HIPE (Ott 2010), using the scripts disseminated at the Herschel data reduction workshop in January, 2010.

2.1. Spectroscopy

Spectroscopic observations with the PACS IFU were carried out in chop-nod mode in order to remove the telescope emission. The PACS IFU is a 5 × 5 filled array, with 9.4″ × 9.4″ spaxels. Observations were carried out such that the central IFU spaxel was centered on the target in both nods. Our SDP targets were carried out with a 2″ dither. Each PacsLineSpec observation includes simultaneous observations from 62.93–64.43 and 180.76–190.29 μm. Our PacsRangeSpec observations included, in series, simultaneous observations from 71.81–73.28 and 143.59–146.53 μm, 78.37–79.76 and 156.70–159.47 μm, and 89.28–90.48 and 178.51–180.96 μm. Typical continuum rms for all regions are \(~0.3\) Jy, with 1σ line flux uncertainties of 2–5 × 10^{-18} W/m². Spectral resolution ranges from R ~ 1100–3400. Full details of the observations are given in Dent et al. (in prep).

For our point source flux measurements, we extracted the A and B nods for the central IFU spaxel. For each nod, we then Nyquist binned in wavelength with non-overlapping bins half the width of the instrumental resolution. The mean of the two nods restores the flux of the objects. We apply wavelength dependent flux and aperture corrections released by the PACS development team², and assume a flux calibration uncertainty of 40%.

To extract line fluxes and measure continuum flux from our reduced, nod-combined, flux corrected spectra, we carried out an inverse-error weighted fit of a gaussian and first-order polynomial to the spectra. Using the polynomial fit to each spectral region, we estimated the continuum emission at the rest wavelength for each line. To estimate the error on the continuum, we calculated the error of the mean for the residual of the polynomial fit in a region from 2 to 10 instrumental FWHM from the line rest wavelength. For emission lines, we report the flux of detected lines as the integrated flux of the gaussian line fit. We calculate 1σ fluxes as the integral of a gaussian with height equal to the continuum RMS, and width equal to the instrumental FWHM; we report 3σ upper limits for non-detections.

The [OI] 63 μm emission line is clearly detected towards HD169142, RECX 15, and TW Hya from our SDP sample; with FWHM ~ 90 km s^{-1}. HD169142 and TW Hya are spectrally unresolved. The emission from RECX 15, on the other hand, appears much broader, with a FWHM ~ 150 km s^{-1}; this is a marginal result and could be due to off-center or low S/N observations in the two nods (further discussion in Woitke et al., in prep.). We show our baseline subtracted [OI] 63 μm spectra in Fig. 1. Examination of spectra from the off-center data cubes show no indications of extended or field emission. In Table 1, we show the [OI] 63, [OI] 145, and [CII] 157 fluxes and 3σ upper limits, as well as the measured continuum fluxes and 3σ upper limits. At this stage, no other lines were detected with high significance.

2.2. Photometry

Simultaneous 70 and 160 μm continuum observations were obtained with the PACS photometer. The blue camera consists of a 64×32 array of 3.2″ pixels, while the red camera consists of a 32×16 array of 6.4″ pixels. Point source mode consists of 50″ chord-nod observations for sky subtraction and a 3 point dither (8.5″) for higher spatial sampling. Scan maps were carried out with a 20″/s scan rate, with a 10 Hz integration cycle.

Footnotes:

1. https://nhscsci.ipac.caltech.edu/sc/index.php/Workshops/
2. PacsSpectroscopyPerformanceAndCalibration_v1_1_0.pdf

Fig. 1. Baseline subtracted, aperture and flux corrected [OI] 63 μm spectra. Error bars indicate 1σ statistical errors; gaussian line fits are shown in red.
Table 1. Stellar properties and observational results.

<table>
<thead>
<tr>
<th>Object</th>
<th>SpT</th>
<th>Age Myr</th>
<th>Distance pc</th>
<th>[OI] 63(^a) (10^{-18}) W/m(^2) Jy</th>
<th>[OI] 145(^a) (10^{-18}) W/m(^2) Jy</th>
<th>[CI] 157(^a) (10^{-18}) W/m(^2) Jy</th>
<th>70(^b) (10^{-18}) W/m(^2) Jy</th>
<th>160(^b) (10^{-18}) W/m(^2) Jy</th>
</tr>
</thead>
<tbody>
<tr>
<td>HD 169142</td>
<td>A5(^{11})</td>
<td>6(^{12})</td>
<td>14.5(^{11})</td>
<td>71.7 ± 3.8</td>
<td>18.89 &lt; 10.4</td>
<td>13.76 &lt; 6.4</td>
<td>14.49</td>
<td>27.35 ± 0.03</td>
</tr>
<tr>
<td>HD 181327</td>
<td>F5(^{d})</td>
<td>12(^{8})</td>
<td>50.6(^{d})</td>
<td>&lt;9.3</td>
<td>1.75 &lt; 8.0</td>
<td>0.80 &lt; 8.5</td>
<td>0.77</td>
<td>1.90 ± 0.03</td>
</tr>
<tr>
<td>TW Hya</td>
<td>K7(^{e})</td>
<td>10(^{20})</td>
<td>56(^{g})</td>
<td>36.5 ± 3.6</td>
<td>2.99 &lt; 5.5</td>
<td>7.00 &lt; 6.6</td>
<td>8.79</td>
<td>3.90 ± 0.02</td>
</tr>
<tr>
<td>RECX 15</td>
<td>M2(^{e})</td>
<td>9(^{21})</td>
<td>97(^{12})</td>
<td>31.5 ± 3.1</td>
<td>0.14 &lt; 4.6</td>
<td>0.10 &lt; 7.4</td>
<td>&lt;0.09</td>
<td>0.18 ± 0.02</td>
</tr>
</tbody>
</table>

Notes. \(^{a}\) Calibration uncertainty 40%; \(^{b}\) 1σ ± 0.03–0.05 Jy; \(^{c}\) 1σ ± 0.02–0.04 Jy; \(^{d}\) 1σ ± 0.03–0.05 Jy; \(^{e}\) Calibration uncertainty 10%; \(^{f}\) Results from two stacked scan maps.

References. (1) Dunkin et al. (1997); (2) Grady et al. (2007); (3) van Boekel et al. (2005); (4) Nordstrom et al. (2004); (5) Zuckerman & Song (2004); (6) Perryman et al. (1997); (7) de la Reza et al. (1989); (8) Barrado Y Navascués (2006); (9) Wichmann et al. (1998); (10, 11) Lawson et al. (2002); (12) Mamajek et al. (2000)

All sources appear spatially unresolved, with \(FWHM \sim 5''\) and \(11''\) at 70 and 160\(\mu\)m, respectively. We carried out photometry on images at the instrumental pixel size using a 21'' aperture for both blue and red camera observations. The PACS team has reported flux corrections and aperture corrections, resulting in a flux calibration uncertainty of 5% in the blue camera, and 10% in the red camera. All four of our SDP targets are detected in both bands; we report these continuum measurements without color correction in Table 1. None of the sources show signs of extended emission. Comparison with publicly available IRAS and Spitzer photometry shows a general agreement within the systematic uncertainties. Our model fitting efforts of lines and continuum from HD 169142 and TW Hya, using MCFOST and ProDiMo, are presented in Meeus et al. (2010) and Thi et al. (2010); RECX 15 and HD 181327 are addressed in further detail in Woitke et al. (in prep) and Ménard et al. (in prep).

We find one exception to a good agreement between continuum measurements from spectroscopy and photometry – blue camera observations of HD 169142 show differences up to 50%. The cause of the difference is currently unknown, and may be due to large systematic flux uncertainties. We will revisit this issue with future revisions of the instrumental calibration and reduction techniques. As the 70\(\mu\)m flux from the photometric observation is more consistent with photometry in the literature, we consider it the more reliable of the two flux values, and note the [OI] line flux for HD 169142 may be higher by up to 50%.

3. Discussion

3.1. Discs or outflows?

A primary driver of the GASPS project is to use far-infrared emission lines to characterize the gas content of circumstellar discs. However, many young star with accreting discs are also known to drive outflows. Their shocks produce strong line emission in this band (Liseau et al. 2006). Following Hollenbach (1985), we can bound the mass outflow rate from the [OI] 63 luminosity, \(M_{\text{outflow}}(M_\odot/\text{yr}) \approx 10^{-4} L_{\text{[O I]63}}(L_\odot)\). As some shock emission may emerge through other lines, for the observed fluxes, this translates into lower limits of 46, 3.5, and 9.1 \(10^{-16} M_\odot/\text{yr}\) for HD 169142, TW Hya, and RECX 15 respectively. Comparison to observed accretion rates (Sec. 1) rule out an outflow shock origin for the emission from HD 169142, and requires the outflow rate to be a very high fraction (≥0.2) of the accretion rate for TW Hya and RECX 15. This is much greater than the ratio found in T Tauri stars by Hartigan et al. (1995).

3.2. [CII] non-detections

The lack of detections of the [CII] 158 line from any of our SDP targets is surprising. Kamp et al. (2010) predict [CII] 158 emission of \(10^{-17} \text{W/m}^2\) for HAe discs with radii \(\sim 200\text{AU}, \)
with emission varying with disc surface area. HD 169142 is known from mm imaging to have a disc of radius >200 AU (Raman et al. 2006). Other factors must be at play, such as flat discs, lack of UV flux, or shielding by a large amount of small dust grains at large radii.

3.3. A new regime of disc studies

The ratio of [O I] 63 line strength to continuum varies substantially, from 0.3–1 for HD 169142 and TW Hya to 17 for RECX 15. Both line and continuum are likely optically thick and the stellar luminosities are very different for the three sources but one possible explanation for the strikingly high line-to-continuum ratio in RECX 15 is an enhanced gas-to-dust ratio (see Fig. 3 of Pinte et al. 2010).

Further, whereas HD 169142 and TW Hya have strong mm continuum fluxes indicating large dust masses, RECX 15 is undetected at 850 μm (Phillips, in prep.). This is similar to the 49 Ceti disc which is detected at 1.3 mm in CO 2–1 but not in the continuum (Hughes et al. 2008). These comparisons suggest a high gas-to-dust ratio; however, further modeling is necessary (see Woitke et al., in prep.). The GASPS survey will reveal how common are discs with gas emission yet lacking detectable dust.

Prior studies of circumstellar disc gas have focused on near-infrared emission lines from the hot inner disc (e.g. Brittain et al. 2007) or mm lines from the cool outer disc (e.g. Dent et al. 2005). Far-infrared line emission from intermediate radii of gas rich discs has not yet been explored in similar detail. Thanks to Herschel’s unprecedented sensitivity, GASPS will examine the gas content of circumstellar discs in previously unaccessible regimes.

Acknowledgements. We thank the PACS instrument team for their dedicated support. G.S.M., D.R.A., S.D.B., W.D., C.A.G., C.D.H., L.P., B.R., A.R., G.S. and J.P.W. acknowledge NASA/JPL for funding support. C.E., G.M., L.M., and B.M. are partly supported by Spanish grant AYA 2008-01727, and J.M.A. and E.S. by Spanish grant AYA2008-02156. C.P. acknowledges the funding from the EC 7th Framework Program as a Marie Curie Intra-European Fellow (PIEF-GA-2008-220891). E.S. and J.M.A. acknowledge funding from the Spanish MICINN through grant AYA2008-02156. The members of LAOG acknowledge PNPS, CNES and ANR (contract ANR-07-BLAN-0221) for financial support. This work is based on observations made with Herschel, a European Space Agency Cornerstone Mission with significant participation by NASA. Support for this work was provided by NASA through an award issued by JPL/Caltech. PACS has been developed by a consortium of institutes led by MPE (Germany) and including UVIE (Austria); KU Leuven, CSL, IMEC (Belgium); CEA, LAM (France); MPIA (Germany); INAF-IASF/OA/OAP/INAF-LENS, SISSA (Italy); IAC (Spain). This development has been supported by the funding agencies BMVIT (Austria), ESA-PRODEX (Belgium), CEA/CNES (France), DLR (Germany), ASI/INAF (Italy), and CIICYT/MCYT (Spain). HIPE is a joint development by the Herschel Science Ground Segment Consortium, consisting of ESA, the NASA Herschel Science Center, and the HIFI, PACS and SPIRE consortia.

References

Hollenbach, D. 1985, Icarus, 61, 36
Mizumoto, K.-I. Morita, & M. Ohishi, ASP Conf. Ser., in press

Page 5 is available in the electronic edition of the journal at http://www.aanda.org