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# The impact of primordial supersonic flows on early structure formation, reionization and the lowest-mass dwarf galaxies

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## ABSTRACT

Tselikhovich and Hirata recently discovered that higher order corrections to the cosmological linear-perturbation theory lead to supersonic coherent baryonic flows just after recombination (i.e.  $z \approx 1020$ ), with rms velocities of  $\sim 30 \text{ km s}^{-1}$  relative to the underlying dark matter distribution, on comoving scales of  $\lesssim 3 \text{ Mpc } h^{-1}$ . To study the impact of these coherent flows, we performed high-resolution  $N$ -body plus smoothed particle hydrodynamic simulations in boxes of 5.0 and 0.7  $\text{Mpc } h^{-1}$ , for bulk-flow velocities of 0 (as reference), 30 and 60  $\text{km s}^{-1}$ . The simulations follow the evolution of cosmic structures by taking into account detailed, primordial, non-equilibrium gas chemistry (i.e. H, He, H<sub>2</sub>, HD, HeH, etc.), cooling, star formation and feedback effects from stellar evolution. We find that these bulk flows suppress star formation in low-mass haloes (i.e.  $M_{\text{vir}} \lesssim 10^8 M_{\odot}$  until  $z \sim 13$ ), lower the abundance of the first objects by  $\sim 1$ –20 per cent and as a consequence delay cosmic star formation history by  $\sim 2 \times 10^7 \text{ yr}$ . The gas fractions in individual objects can change by up to a factor of 2 at very early times. Coherent bulk flow therefore has implications for (i) the star formation in the lowest-mass haloes (e.g. dSphs); (ii) the start of reionization by suppressing it in some patches of the Universe; and (iii) the heating (i.e. spin temperature) of neutral hydrogen. We speculate that the patchy nature of reionization and heating on several Mpc scales could lead to enhanced differences in the H I spin temperature, giving rise to stronger variations in the H I brightness temperatures during the late dark ages.

**Key words:** galaxies: evolution – galaxies: formation – cosmology: theory – dark ages, reionization, first stars – early Universe.

## 1 INTRODUCTION

Cosmic structure-formation models are based on Jeans theory (Jeans 1902) applied to primordial matter fluctuations in an expanding Universe (e.g. Komatsu et al. 2010). The matter density perturbations are linearly expanded to first order, which is supposed to work well as long as the density contrast is much smaller than unity (e.g. Peebles 1974; Coles & Lucchin 2002; Ciardi & Ferrara 2005), and the equations that govern the evolution of the dark and baryonic matter become linear. To study highly non-linear evolution when the density contrast approaches and far exceeds unity, however, fully numerical simulations are required. According to the concordance cosmological model (e.g. Komatsu et al. 2010), the Universe is flat and has an expansion rate of  $H_0 \simeq 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ . The cosmological constant (or dark energy) represents the main energy density content,  $\Omega_{0,\Lambda} \simeq 0.7$ , whereas dark matter contributes  $\Omega_{0,M} \simeq 0.26$

and baryonic matter contributes only a small fraction,  $\Omega_{0,b} \simeq 0.04$ . Within this context, baryonic structures arise from infall and condensation of gas into growing dark matter potential wells (e.g. Gunn & Gott 1972; Peebles 1974; White & Rees 1978). Both infall and condensation of gas take place only *after* the time of recombination ( $z \simeq 1020$ ), when baryons and photons decouple (Komatsu et al. 2010). On the other hand, dark matter perturbations can already start growing when their density exceeds that of radiation (i.e. at  $z_{\text{eq}} \sim 3100$ –3200). Recently, (Tselikhovich & Hirata 2010) for the first time included quadratic terms in the cosmological perturbation theory to account for the advection of small-scale perturbations by large-scale velocity flows. They find that, at the redshift of decoupling  $z \simeq 1020$ , coherent supersonic flows of the baryons relative to the underlying dark matter distribution are formed on scales of a few Mpc or less, with typical velocities on those scales with a rms value of  $\sim 30 \text{ km s}^{-1}$ , sourced by density perturbations up to scales of  $\sim 100 \text{ Mpc}$ . These gas bulk velocities could actually suppress the formation of the very early and low-mass structures, induce higher baryon acoustic oscillations (Dalal, Pen & Seljak 2010), and

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have further implications on high-redshift galaxy clustering, 21-cm studies and reionization. In fact, as we find, they might impact the formation of the lowest-mass galaxies (e.g. dSphs) on scales as large as the Local Group, depending on the magnitude of the bulk flow at decoupling. The effect can thus cause a spatially varying bias in baryonic structure formation, whereas small-scale baryonic structures are suppressed more in regions where the bulk velocity is larger. Detailed numerical simulations, considering this non-linear effect, are not yet available and the very high dynamic range required is not foreseen in the near future. It therefore looks very challenging to quantitatively assess these effects in the highly non-linear regime discussed by Tselikhovich & Hirata (2010). Indeed, one would need to resolve, at the same time, large scales (of the order of the horizon) and extremely small scales (of the order of the Jeans length).

Given these extreme difficulties, we adopt the following approach: we focus on small scales where the effect is expected to be largest, at least properly dealing with all the detailed physics and chemistry, and we assume a *constant bulk velocity shift in the initial conditions* as determined by Tselikhovich & Hirata (2010) on the relevant scales. This is a simple, but reasonable assumption to study the lowest-mass structures in the box (which have virial radii much smaller than the box size) and we expect it to be able to provide the first quantitative results in the highly non-linear regime and assess its effects on early structure formation. In this Letter, we take the first step towards understanding the effects of coherent supersonic bulk flow on the formation of the smallest baryonic structures, in particular focusing on the suppression of star formation and the delay in reionization. In Section 2 we describe our simulations and in Section 3 we show our results. We summarize and conclude in Section 4.

## 2 SIMULATIONS

We perform numerical simulations by using the parallel tree/smoothed particle hydrodynamics (SPH) P-GADGET2 code (Springel 2005), which implements gravity, hydrodynamics, molecular and atomic evolution, cooling, Population II and Population III star formation, supernova/wind feedback, and a full chemistry network involving  $e^-$ , H,  $H^+$ ,  $H^-$ , He,  $He^+$ ,  $He^{++}$ ,  $H_2$ ,  $H_2^+$ , D,  $D^+$ , HD and  $HeH^+$  (for more details see Maio et al. 2006, 2007, 2010a; Maio 2009; Maio et al. 2010b). Particularly important for our purposes is the capability of the code of following gas collapse down to the catastrophic molecular cooling branch by resolving the Jeans scales of primordial structures with  $\sim 10^2$  particles (Maio et al. 2009) and completely fulfilling the correct resolution requirements for SPH simulations (see Bate & Burkert 1997; Maio et al. 2009, for details). We set up the initial conditions by using the N-GENIC code and add bulk velocity shifts for the gas in the  $x$ -direction  $v_{b,x} = 0, 30$  and  $60 \text{ km s}^{-1}$  (referred to as HR-Shift00, HR-Shift30 and HR-Shift60, respectively) at recombination epoch and properly scaled down (according to Tselikhovich & Hirata 2010) to the simulation initial redshift of  $z_{in} = 100$ . The value  $v_{b,x} \sim 60 \text{ km s}^{-1}$  has been considered to account for statistical variations from the expected rms value of  $\sim 30 \text{ km s}^{-1}$ , but it should be considered as an upper limit, which can be found in less than 1 per cent of the volume of the Universe. Even though we expect that the presence of bulk flows at higher redshift does not affect the structure formation process, we nevertheless run additional simulations with  $z_{in} = 1020$ , with  $v_{b,x} = 0$  and  $60 \text{ km s}^{-1}$  (HR-Shift00Rec and HR-Shift60Rec, respectively). We should caution the reader that the P-GADGET2 code implements  $N$ -body/SPH structure formation, cooling, star formation and feedback effects in

matter-dominated universes, while contributions from radiation or from the primordial hot plasma in the cosmological dynamics are neglected. Therefore, its use at  $z \gg 200$  might be questionable. In any case, the results are very similar to the simulations with  $z_{in} = 100$  and thus we do not discuss them any further. Our reference simulations have a box side of  $0.7 \text{ Mpc } h^{-1} \simeq 1 \text{ Mpc}$  and are run in the frame of the standard  $\Lambda$  cold dark matter cosmological model ( $\Omega_{0,\Lambda} = 0.7$ ,  $\Omega_{0,M} = 0.26$ ,  $\Omega_{0,b} = 0.04$ ,  $h = 0.7$ ,  $\sigma_8 = 0.9$ ,  $n = 1$ ). Matter and gas fields are sampled with  $320^3$  particles, respectively, allowing a resolution of  $\sim 10^2 \text{ M}_\odot h^{-1}$  for the gas component and  $\sim 7.5 \times 10^2 \text{ M}_\odot h^{-1}$  for the dark-matter component. The chemical abundances are initialized as in Maio et al. (2007, 2010a). To check if the effects of primordial supersonic gas flows on structure formation could be seen in bigger but lower resolution simulations, we also run a set of three larger boxes (LR-Shift00, LR-Shift30 and LR-Shift60) with a size of  $5 \text{ Mpc } h^{-1}$ ,  $v_{b,x} = 0, 30$  and  $60 \text{ km s}^{-1}$ , respectively, and with the same parameters as described before. This box is comparable to the smallest scale where the bulk flows are expected according to Tselikhovich & Hirata (2010). The mass resolutions for gas and dark matter species are  $\sim 4 \times 10^4$  and  $\sim 3 \times 10^5 \text{ M}_\odot h^{-1}$ , respectively. In this case, we did not find any relevant differences between the three runs, implying that the effect is observable *only* if the numerical resolution is high enough to resolve halo masses of  $\sim 10^4 \text{ M}_\odot h^{-1}$ . This highlights the difficulty of present large-scale simulations to address these effects. In Table 1, we list the main features of the numerical set-ups.

## 3 RESULTS

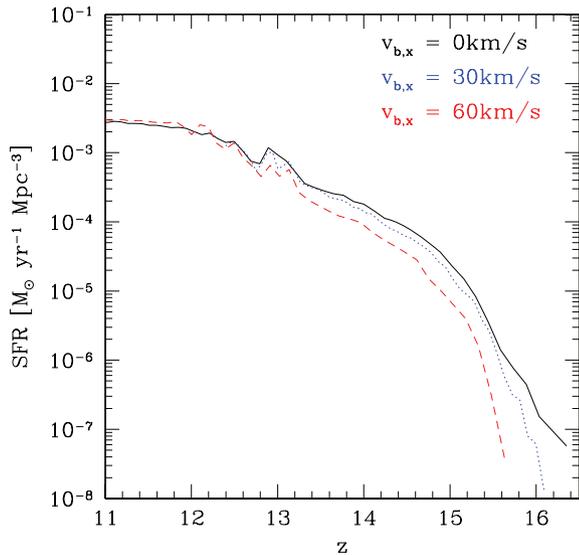
We start by looking at the overall evolution of the star formation rate density of the Universe and then we consider the main statistical quantities involved in the cosmological evolution process, that is, the distributions of dark matter haloes and gas clouds.

### 3.1 Star formation rate density

The behaviour of the star formation rate densities is plotted in Fig. 1, for all the cases considered before. What emerges is a systematic shift in the onset of star formation, with a delay increasing for higher values of  $v_{b,x}$ . For the  $v_{b,x} = 0 \text{ km s}^{-1}$  case, the onset happens at  $z \sim 16.3$ , when the Universe is roughly  $2.5 \times 10^8 \text{ yr}$  old, while in the  $v_{b,x} = 30$  and  $60 \text{ km s}^{-1}$  cases, star formation sets in at  $z \sim 16.1$  and  $15.6$ , respectively, corresponding to a delay of  $\sim 10^7$  and  $\sim 2 \times 10^7 \text{ yr}$  compared to the case without velocity shift. Differences at very early times are almost one order of magnitude and at later times a factor of a few, persisting down to  $z \sim 13$ , when the global trends rapidly converge to similar values. The reason behind this is the higher kinetic energy given to the gas at the recombination epoch: because of that, first, dark matter haloes can retain less material, gas condenses more slowly hindering molecule formation, the corresponding cooling is less efficient and the resulting star formation is delayed. In particular, this is very relevant for small primordial haloes, with masses of  $\sim 10^4$ – $10^8 \text{ M}_\odot$  (see also Tselikhovich & Hirata 2010), whose dimensions are comparable to or smaller than the baryon Jeans length. In this case, in fact, the gas cannot fragment within the halo and partially or entirely flows out of it, since larger dark matter potential wells are needed to retain the gas. Therefore, this effect on small scales influences the growth of larger objects as well. From this we conclude that the inflow and condensation of gas, and thus star formation, is already strongly suppressed in haloes in the mass range  $\lesssim 10^8 \text{ M}_\odot$  *before reionization*. Those low-mass haloes surviving to the present day, or even those that merge

**Table 1.** Parameters adopted for the simulations.

Model	Box side (Mpc $h^{-1}$ )	Number of particles	Mean interparticle separation (kpc $h^{-1}$ )	$M_{\text{gas}}$ ( $M_{\odot} h^{-1}$ )	$M_{\text{dm}}$ ( $M_{\odot} h^{-1}$ )	Initial redshift	Bulk shift at $z = 1020$ (km $s^{-1}$ )
HR-Shift00	0.7	$2 \times 320^3$	2.187	$1.16 \times 10^2$	$7.55 \times 10^2$	100	0
HR-Shift30	0.7	$2 \times 320^3$	2.187	$1.16 \times 10^2$	$7.55 \times 10^2$	100	30
HR-Shift60	0.7	$2 \times 320^3$	2.187	$1.16 \times 10^2$	$7.55 \times 10^2$	100	60
HR-Shift00Rec	0.7	$2 \times 320^3$	2.187	$1.16 \times 10^2$	$7.55 \times 10^2$	1020	0
HR-Shift60Rec	0.7	$2 \times 320^3$	2.187	$1.16 \times 10^2$	$7.55 \times 10^2$	1020	60
LR-Shift00	5.0	$2 \times 320^3$	15.62	$4.23 \times 10^4$	$2.75 \times 10^5$	100	0
LR-Shift30	5.0	$2 \times 320^3$	15.62	$4.23 \times 10^4$	$2.75 \times 10^5$	100	30
LR-Shift60	5.0	$2 \times 320^3$	15.62	$4.23 \times 10^4$	$2.75 \times 10^5$	100	60


**Figure 1.** Total star formation rates for bulk velocity shifts of  $0 \text{ km s}^{-1}$  (black solid line),  $30 \text{ km s}^{-1}$  along the  $x$ -axis (blue dotted line) and  $60 \text{ km s}^{-1}$  along the  $x$ -axis (red dashed line) are shown for the simulation box of side  $\sim 1 \text{ Mpc}$ .

into larger masses, would therefore have a lower baryonic and stellar mass fraction. We further note that the effect of suppressing star formation is similar to that of reionization, but it has a very different physical origin and occurs well before reionization itself.

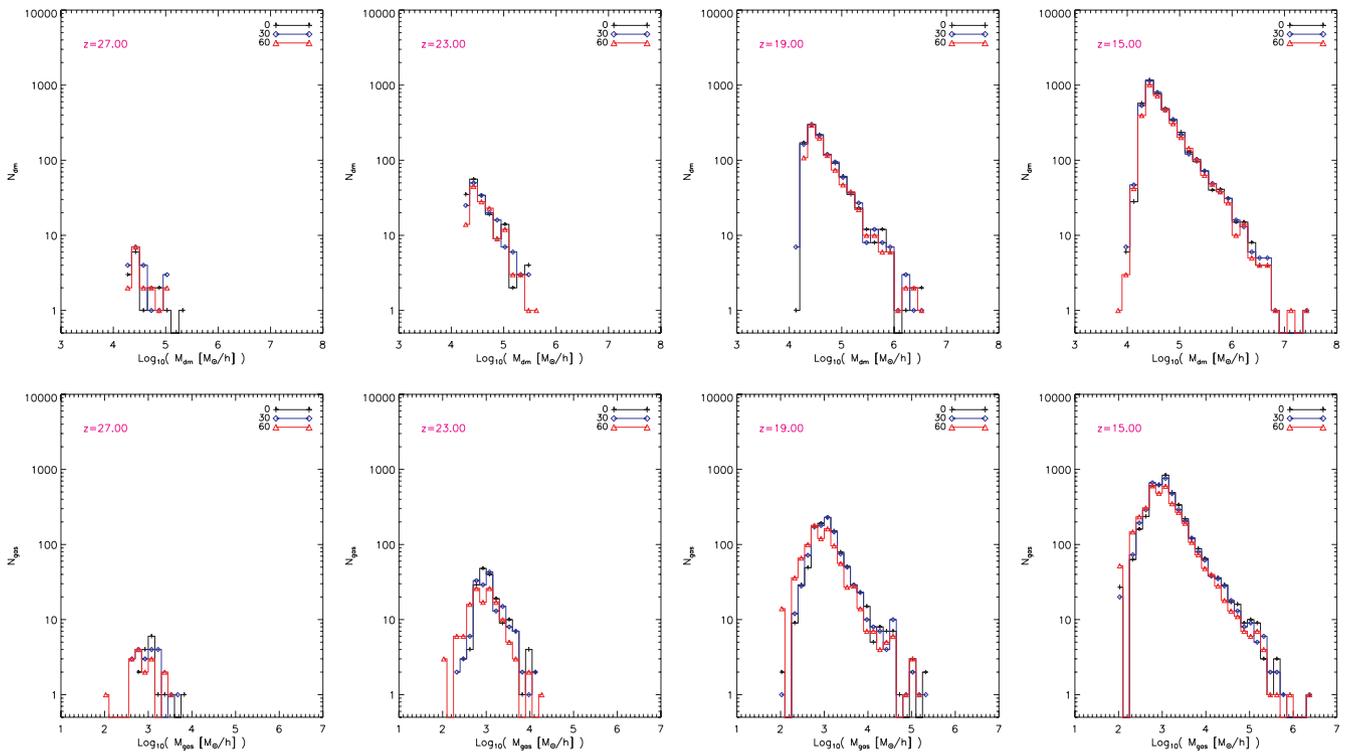
### 3.2 Dark matter haloes and gas clouds

In each simulation, we identify cosmic structures with the use of a friend-of-friend algorithm. Each object is represented by all those particles (dark matter, gas and star) closer than a limiting linking length of 20 per cent the mean interparticle separation. The dark matter halo and gas cloud distributions are shown in Fig. 2, at different redshifts as indicated by the legends, for the runs starting at  $z_{\text{in}} = 100$ . In general, at very high redshift, we cannot give solid conclusions, since small-number statistics dominate the sample. At later times (i.e.  $z \lesssim 23$ ), some systematic behaviour clearly emerges. The mass distributions of dark matter haloes range between  $\sim 10^4$  and  $10^7 M_{\odot} h^{-1}$ . The three cases of  $v_{b,x} = 0, 30$  and  $60 \text{ km s}^{-1}$  show a slight decrement with increasing  $v_{b,x}$ , but only at a few per cent level. Hence, as expected, the effect of the bulk flow of gas on the overall growth of the dark-matter-dominated haloes is minimal.

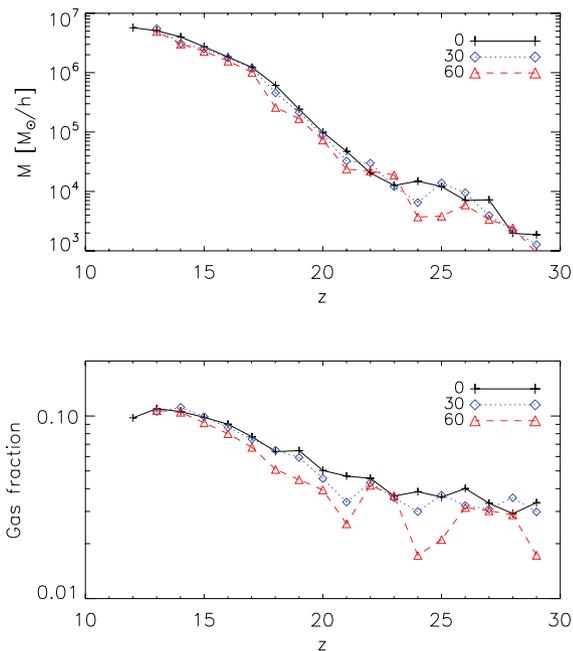
More clear differences arise from the analyses of the gas clouds. In Fig. 2, we also show the mass distribution of the gaseous com-

ponent of the objects and we clearly see decrements up to tens of per cent, mostly evident in the  $v_{b,x} = 60 \text{ km s}^{-1}$  case. Interestingly, there is a cascading effect from smaller to larger masses (expected from the simple analyses of the star formation rates), which delays and slightly hinders the accumulation of gas in the dark matter haloes (i.e. suppression of gas inflow in lower mass haloes also suppresses the gas accumulation in high-mass haloes through accretion of low-mass haloes). As expected, the average gas cloud mass decreases for larger  $v_{b,x}$ . At  $z \simeq 23$ , it drops from  $\sim 1.70 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 0 \text{ km s}^{-1}$ ) down to  $\sim 1.62 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 30 \text{ km s}^{-1}$ ) and  $\sim 1.39 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 60 \text{ km s}^{-1}$ ), while at  $z \simeq 15$ , it is  $\sim 4.41 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 0 \text{ km s}^{-1}$ ),  $\sim 4.26 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 30 \text{ km s}^{-1}$ ) and  $\sim 4.06 \times 10^3 M_{\odot} h^{-1}$  ( $v_{b,x} = 60 \text{ km s}^{-1}$ ). The differences correspond to decrements with respect to the reference run of  $\sim 6\text{--}3$  ( $v_{b,x} = 30 \text{ km s}^{-1}$ ) and  $\sim 20\text{--}10$  per cent ( $v_{b,x} = 60 \text{ km s}^{-1}$ ) at  $z \simeq 23\text{--}15$ . Also the total abundance of gas clouds shows a similar trend, with mass-weighted number counts dropping, with respect to the reference run, by  $\sim 5\text{--}3$  ( $v_{b,x} = 30 \text{ km s}^{-1}$ ) and  $\sim 20\text{--}50$  per cent ( $v_{b,x} = 60 \text{ km s}^{-1}$ ) at  $z \simeq 23\text{--}15$ . We have verified that the excess small clumps observed in the low-mass tail of the distributions (at masses  $\lesssim 500 M_{\odot} h^{-1}$ ) for  $v_{b,x} > 0$  are due to residual gas which, differently from the reference run, is not trapped by the dark matter haloes and later collapses to form such clumps. It should be noted though that statistical errors in this mass regime are substantial due to the small number of particles contained in a halo.

As a more specific example, in Fig. 3, we show the evolution of the gas content (top panel) and gas fraction (bottom panel) of the first (and most-massive) halo formed in the simulations. The gas mass ranges from  $\sim 10^3$  to  $10^7 M_{\odot} h^{-1}$  and the corresponding dark matter mass is about one order of magnitude larger. The gas mass variations due to the different values of  $v_{b,x}$  are significant at  $z \gtrsim 16$ , while at smaller redshift, the total amount converges to  $\sim (2\text{--}6) \times 10^6 M_{\odot} h^{-1}$ . In this respect, the gas fractions in the halo are more informative and suggest a depletion of gas for higher  $v_{b,x}$  up to a factor of 2. Indeed, at  $z \gtrsim 24$ , the gas fraction for the  $v_{b,x} = 0 \text{ km s}^{-1}$  case is  $\sim 0.04$ , while in the  $v_{b,x} = 30$  and  $60 \text{ km s}^{-1}$  cases, it drops to  $\sim 0.03$  and below 0.02, respectively. Similar strong variations are also observed at later times (e.g. at  $z \sim 23$ ), but the general trend converges at  $z \lesssim 16$ , when the discrepancies among the cases become  $< 10$  per cent and all the gas fractions reach values of  $\sim 0.10$ . The same effects are observed in the gas mass and density profiles as a function of the (physical) radius – see Fig. 4 and compare to fig. 2 in Tseliakhovich & Hirata (2010). At early times (left-hand panels), the amount of gas falling in the halo depends on  $v_{b,x}$ , the larger  $v_{b,x}$ , the more gas depletion in the distribution tails. At larger radii, the mass depletion is up to  $\sim 50$  per cent and smoothly decreases towards the core, where the gas is even denser and more



**Figure 2.** Top row: dark matter halo distributions at different redshifts (see labels) for the 1-Mpc-side boxes. Bottom row: gas cloud distributions at different redshifts (see labels) for the 1-Mpc-side boxes. Histograms refer to velocity shifts of  $0 \text{ km s}^{-1}$  (black crosses),  $30 \text{ km s}^{-1}$  along the  $x$ -axis (blue rhombi) and  $60 \text{ km s}^{-1}$  along the  $x$ -axis (red triangles).



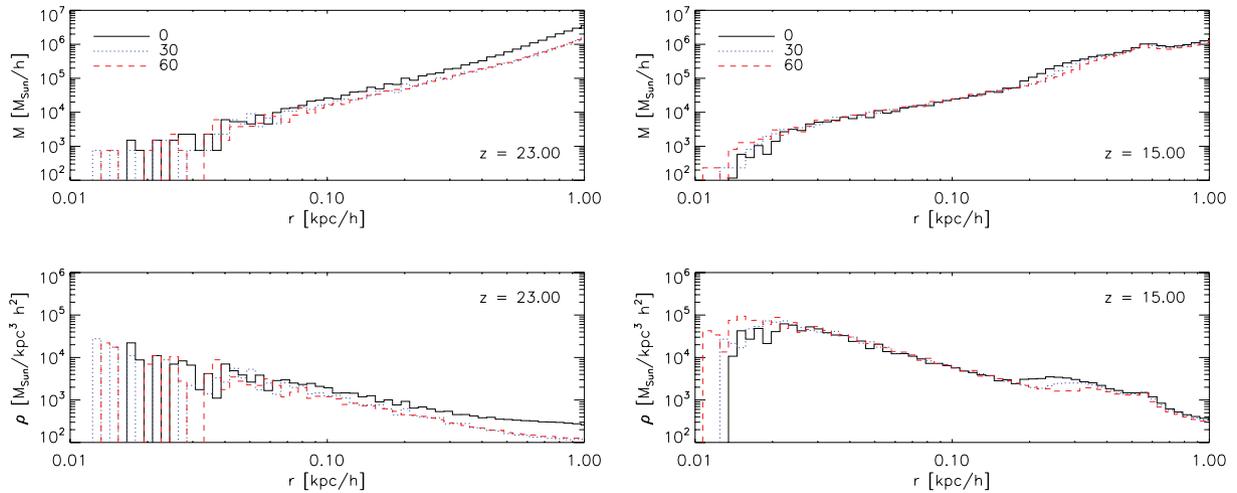
**Figure 3.** Evolution of the gas content (top panel) and of the gas fraction (bottom panel) for the first (and most-massive) halo formed in the simulations. Data refer to velocity shifts of  $0 \text{ km s}^{-1}$  (solid lines, black crosses),  $30 \text{ km s}^{-1}$  (dotted lines, blue rhombi) and  $60 \text{ km s}^{-1}$  (dashed lines, red triangles).

abundant. This can be seen as a consequence of the fact that there is less gas falling in and it sinks more easily in the potential wells of the halo. Of course, given the low number of SPH particles reaching those regimes, the statistical errors might be quite significant

and ad hoc simulations are needed to address this issue in more detail.

#### 4 SUMMARY AND CONCLUSIONS

Stimulated by very recent analytical work (Tselikhovich & Hirata 2010), we have run cosmological, high-resolution numerical  $N$ -body/SPH simulations including detailed, non-equilibrium chemistry evolution, cooling, star formation and feedback effects (see e.g. Maio et al. 2007, 2010b), which consistently follow the gas behaviour down to the catastrophic cooling regime, and correctly resolve the small, primordial Jeans scales with  $\sim 10^2$  SPH particles. To take into account very early, supersonic, coherent, Mpc-scale flows, generated at recombination time from the advection of small-scale perturbations by large-scale velocity flows, we set initial gas velocity shifts along the  $x$ -direction,  $v_{b,x} = 0, 30$  and  $60 \text{ km s}^{-1}$ , based on analytic estimates of the rms differences in dark matter and baryonic gas velocities at  $z = 1020$ . We find that in the simulations with larger  $v_{b,x}$ , the onset of star formation is delayed by some  $\sim 10^7$  yr and it is smaller by a factor of a few in comparison to the  $v_{b,x} = 0 \text{ km s}^{-1}$  case, until redshift  $z \sim 13$ . As a consequence, the beginning of the reionization process is delayed by the same amount and initially driven by more massive objects. More ad hoc calculations should be done to quantify the differences in terms of evolution, topology and photon budget. This might also have consequences in terms of detection of the 21-cm signal from neutral hydrogen, because the building-up of a Ly $\alpha$  background (see e.g. Robertson et al. 2010, for a recent review) necessary for the observability of the line (e.g. Furlanetto, Oh & Briggs 2006) would be delayed as well. We speculate that this could lead to more patchy reionization and H I heating, with larger H I brightness-temperature fluctuations. More research is needed to quantify these effects.



**Figure 4.** Gas profiles as a function of the (physical) radius at redshift  $z = 23$  (left-hand panels) and  $z = 15$  (right-hand panels) for the data referring to the simulations with velocity shifts of  $0 \text{ km s}^{-1}$  (solid black lines),  $30 \text{ km s}^{-1}$  (dotted blue lines) and  $60 \text{ km s}^{-1}$  (dashed red lines). The top panels show the gas mass profiles, while the bottom panels show the gas density profiles.

Because we find that the amount of gas in the first haloes is expected to drop significantly (up to a factor of 2 at early times i.e.  $\sim 50$  per cent), there could also be interesting consequences in terms of feedback effects, as, for example, less gas in small-mass objects should make it easier to further remove gas via photoevaporation, winds or supernova explosions. It will be interesting to investigate in more details the interplay between bulk-flow-induced feedback effects. Slight differences (at  $\sim 1$  per cent level) in the dark matter statistical distributions also arise, but stronger differences (up above  $\sim 10$  per cent level) are seen in the gas clouds. A cascade effect transfers the inability of primordial, small haloes to completely retain the gas to larger, lower redshift haloes whose gaseous components are systematically smaller for  $v_{b,x} > 0 \text{ km s}^{-1}$ . In this way, the original hindering for the very first objects is transferred to the structure which forms later on, affecting their growth history. Numerical studies performed on lower resolution simulations show no detectable differences among the different cases and this means that the lack of numerical resolution can be a strong limitation to probe higher order corrections to linear perturbation theory and SPH particle masses of  $\sim 10^2 M_{\odot} h^{-1}$  are required. The precise initial redshift does not alter our findings significantly either.

Our general conclusions from this albeit preliminary, but *first* fully non-linear, study of the effects of bulk gas flows on non-linear cosmic structure growth are that it: (i) delays early star formation, which affects both reionization (by delaying it to lower redshifts) and the heating of gas of higher redshift, possibly affecting total/global emission and absorption features of H I against the cosmic microwave background by making these effects more patchy; (ii) suppresses star formation in the lowest-mass haloes over all redshifts, which has a cascading effect over nearly 300 Myr on higher mass haloes, which are their merger products. Because these small early objects, if they survive till the present day, are expected to be the parent population of, for example, dSph galaxies, the effect of the bulk flows can also have an impact on the stellar/gas content of the lowest-mass dwarf satellites around more massive galaxies (e.g. the Milky Way) and be highly spatially dependent (i.e. the effect is only present where there were large bulk flows at decoupling). This could result in a strongly spatially varying bias as suggested by (Tselikhovich & Hirata 2010) and (Dalal et al. 2010). In addition,

suppression of gas infall and condensation on these small scales could play an important role in the ‘missing satellite problem’ (see e.g. Kravtsov 2010, and references therein).

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