ABSTRACT

We report the discovery of a giant H I tail in the intragroup medium of HCG 44 as part of the ATLAS$^{3D}$ survey. The tail is ~300 kpc long in projection and contains $\sim 5 \times 10^8$ M$_\odot$ of H I. We detect no diffuse stellar light at the location of the tail down to $\sim 28.5$ mag arcsec$^{-2}$ in g band. We speculate that the tail might have formed as gas was stripped from the outer regions of NGC 3187 (a member of HCG 44) by the group tidal field. In this case, a simple model indicates that about 1/3 of the galaxy’s H I was stripped during a time interval of $<1$ Gyr. Alternatively, the tail may be the remnant of an interaction between HCG 44 and NGC 3162, a spiral galaxy now $\sim 650$ kpc away from the group. Regardless of the precise formation mechanism, the detected H I tail shows for the first time direct evidence of gas stripping in HCG 44. It also highlights that deep H I observations over a large field are needed to gather a complete census of this kind of events in the local Universe.

Key words: galaxies: evolution – galaxies: groups: general – galaxies: groups: individual: HCG 44 – galaxies: interactions – galaxies: ISM.
1 INTRODUCTION

Deep observations of neutral hydrogen (H\textsc{i}) around galaxies often reveal faint gaseous distributions which do not trace galaxies’ stellar light. Known cases include extremely large H\textsc{i} discs and rings around both late-type galaxies (e.g. Krumm & Burstein 1984; Meurer et al. 1996; Koribalski & López-Sánchez 2009) and early-type galaxies (e.g. van Driel & van Woerden 1991; Morganti et al. 2006; Oosterloo et al. 2007; Serra et al. 2012), tidal tails around merger remnants (e.g. Schiminovich et al. 1994; Koribalski, Gordon & Jones 2003; Duc et al. 2011), tidal and ram-pressure tails in galaxy groups and clusters (e.g. Verheijen & Zwaan 2001; Davies et al. 2004; Oosterloo & van Gorkom 2005; English et al. 2010; Scott et al. 2012; Serra et al. 2012), gas accretion signatures (Sancisi et al. 2008, and references therein) and objects whose origin is still debated (e.g. Schneider et al. 1983; Thilker et al. 2009; Michel-Dansac et al. 2010). These systems carry invaluable information on the way galaxies assemble their stellar mass and accrete and lose gas in different environments.

Neutral hydrogen distributions revealing ongoing tidal interaction, galaxy merging and gas stripping are particularly common in groups of galaxies (e.g. van der Hulst 1979; Verdes-Montenegro et al. 2001; Kern et al. 2008; Freeland, Stilp & Wilcots 2009; see also Hibbard et al. 2001, and references therein). The observation of these processes in action suggests that the morphology and gas content of galaxies can undergo substantial evolution inside a group. In fact, group processes might be a major driver of the morphology–density relation (Dressler 1980; Postman & Geller 1984; for a discussion of the role of groups in determining this relation see e.g. Wilman et al. 2009; Bekki & Couch 2011; Cappellari et al. 2011b) and of the decrease of galaxies’ H\textsc{i} content with increasing environment density (e.g. Verdes-Montenegro et al. 2001; Kilborn et al. 2009; Rasmussen et al. 2012a; Serra et al. 2012).

Strong indications of the importance of group processes come from a combination of optical and H\textsc{i} results presented recently as part of the ATLAS\textsuperscript{3D} study (Cappellari et al. 2011a). First, the fraction of fast rotating early-type galaxies increases (and the fraction of spiral galaxies decreases) with environment density following a surprisingly tight log–linear relation, which is steeper and better defined when the density is measured on a group-like scale (defined by the distance from the third closest neighbour) rather than a cluster-like scale (tenth neighbour; Cappellari et al. 2011b). Secondly, the H\textsc{i} morphology of gas-rich early-type galaxies appears to be strongly related to environment density when the latter is measured on a group-like scale (Serra et al. 2012). Large, regular, settled H\textsc{i} distributions are typical in poor environments (where the distance from the third neighbour is larger than a few Mpc). More disturbed distributions are typical in galaxy groups and may be revealing the processes responsible for the morphology–density relation.

We note that understanding group processes is important to understand galaxy properties over a broad range of large-scale environment densities. On one hand, in a Λ cold dark matter Universe galaxy clusters grow by accretion of groups of (rather than isolated) galaxies, and pre-processing in groups may be important to shape the properties of galaxies living in clusters at redshift zero. On the other hand, even inside large-scale voids galaxies live clustered in small groups and their evolution is to some extent driven by tidal interactions and merging (Szomoru et al. 1996; Kreckel et al. 2012).

Within this context, stripping of gas from galaxies in groups plays a key role. Although direct detection of stripped gas has been possible in a number of groups (see references above), in many other systems galaxies are found to be H\textsc{i} deficient but no intragroup H\textsc{i} is detected. From an observational point of view the main challenge is that the stripped gas can diffuse quickly in the group medium (in about a group crossing time). It therefore reaches low column densities, requiring high sensitivity, and spreads over large areas, requiring observations over a large field. At the moment it is therefore unclear how ubiquitous gas stripping is in groups and, in more quantitative terms, what fraction of the original H\textsc{i} mass of a galaxy can be stripped and on what time-scale. These estimates are needed to understand the importance of gas stripping relative to other processes which may play a role in determining the morphology–density relation, such as the decrease of the cold-gas accretion rate in denser environments.

In this paper, we report the discovery of a giant H\textsc{i} tail in the galaxy group HCG 44 as part of the ATLAS\textsuperscript{3D} survey. We summarize the properties of the group in Section 2, describe radio and optical observations in Section 3, present and discuss the results in Sections 4 and 5, and draw conclusions in Section 6.

2 HICKSON COMPACT GROUP 44

HCG 44 is a compact group at a distance of 25 Mpc (see below) hosting four galaxies of comparable magnitude within an area of 15×15 arcmin\textsuperscript{2} (Fig. 1; Hickson 1982). The two galaxies in the centre of the figure, NGC 3187 and NGC 3190,\textsuperscript{1} exhibit signs of tidal interaction. The former is a blue, late-type spiral with long tails pointing 90\textdegree away from the plane of the galaxy. The latter is an earlier type system with a disturbed morphology and a prominent dust lane. Vorontsov-Velyaminov (1959) groups these galaxies together in the object VV 307 (later catalogued as

\textsuperscript{1} Some authors refer to this galaxy as NGC 3189. According to the NASA Extragalactic Database NGC 3189 is the south-western side of NGC 3190.
Table 1. Galaxies in the HCG 44 field.

<table>
<thead>
<tr>
<th>Galaxy</th>
<th>$v_{hel}$ (km s$^{-1}$)</th>
<th>Distance (Mpc)</th>
<th>Method</th>
<th>Reference</th>
</tr>
</thead>
<tbody>
<tr>
<td>NGC 3185</td>
<td>1230</td>
<td>20 ± 5</td>
<td>TF</td>
<td>a</td>
</tr>
<tr>
<td>NGC 3187</td>
<td>1590</td>
<td>26 ± 10</td>
<td>TE</td>
<td>b</td>
</tr>
<tr>
<td>NGC 3190</td>
<td>1300</td>
<td>24 ± 5</td>
<td>SNei</td>
<td>c, d, f, g, h, i</td>
</tr>
<tr>
<td>NGC 3193</td>
<td>1381</td>
<td>34 ± 3</td>
<td>SBF</td>
<td>j</td>
</tr>
<tr>
<td>SDSS J1017</td>
<td>1940</td>
<td>29</td>
<td>HF</td>
<td>This work</td>
</tr>
</tbody>
</table>

Columns are listed as follows. (1) Galaxy name. (2) Heliocentric velocity measured from the H$\alpha$ data discussed in this paper except for NGC 3193, for which we use the value given by Cappellari et al. (2011a). (3) Redshift-independent distance for all galaxies except SDSS J1017, for which we assume Hubble flow (see below). Error bars are taken from the references in column 5 except for NGC 3190, for which we give the range of distances obtained by different groups who studied the two SN Ia 2002bo and 2002cv; (4) Method used to determine the distance; TF = Tully–Fisher relation corrected for Malmquist bias; TE = Tully estimate; SBF = surface brightness fluctuation corrected for Malmquist bias; HF = Hubble flow with $h=0.73$ after correction for Virgo infall by $+200$ km s$^{-1}$. (5) References for the distance estimate: (a) Springob et al. (2009); (b) Tully (1988); (c) Reindl et al. (2005); (d) Wang et al. (2006); (e) Elias-Rosa et al. (2008); (f) Takanashi, Doi & Yasuda (2008); (g) Wood-Vasey et al. (2008); (h) Mandel et al. (2009); (i) Amanullah et al. (2010); and (j) Blakeslee et al. (2001).

Arp 316), suggesting that they are interacting. The early-type galaxy to the north-east, NGC 3193, is very close in projection to these two galaxies and has regular morphology. To the south-west, NGC 3185 is a barred galaxy with a warped, star-forming outer ring. In addition to these four galaxies listed in the original Hickson (1982) catalogue, the much smaller galaxy SDSS J101723.29+214757.9 (hereafter SDSS J1017) might be at larger distance (see below), while the small blue galaxy in the south-east corner is definitely a background object based on its Sloan Digital Sky Survey (SDSS) redshift of 0.014.

The group membership is not well established. Table 1 lists recessional velocity and distance estimates of the galaxies mentioned above. NGC 3185, NGC 3190 and NGC 3193 have comparable velocities. Relative to these, the velocity of NGC 3187 seems too large for such a small group. SDSS J1017 has even higher velocity, casting doubts on its membership.

Redshift-independent distance estimates of individual galaxies are also puzzling. Taken at face value, the distance of NGC 3193 would put it in the background, as suggested also by Aguerri et al. (2006) based on the non-detection of planetary nebulae in this galaxy. However, error bars on individual distance measurements are large and systematic effects may be important. For example, the two Type Ia supernovae (SNei) in NGC 3190 (see table) are extremely obscured and dust corrections are substantial. Furthermore, NGC 3185 has a warp and this would cause a systematic error on the Tully–Fisher distance based on single-dish H$\alpha$ data.

In this paper, we do not attempt to resolve the issue of galaxy distances and group membership. Instead, we assume that all galaxies labelled in Fig. 1 belong to the group with the exception of SDSS J1017, which we consider as a background object based on its large recessional velocity. We assume that the group is at a distance of $d_{CGAL} = 25$ Mpc. At this distance 15 arcmin correspond to $\sim 110$ kpc.

We note that HCG 44 is part of a loose overdensity which includes 10 to 16 bright galaxies depending on the grouping criterion (see group 58 in Geller & Huchra 1983 and group 194 in Garcia 1993). Galaxies in this overdensity have recessional velocity between $\sim 1100$ and $\sim 1500$ km s$^{-1}$ and are distributed over a sky area of about $5 \times 3$ deg$^2$, which corresponds to $2.2 \times 0.9$ Mpc$^2$ at the distance of HCG 44. Therefore, galaxies in this region are relatively distant from each other (each of them occupies on average 0.6 deg$^2 \sim 0.1$ Mpc$^2$). In comparison, the galaxy number density within HCG 44 is 10 times larger, and this group stands out clearly as a strong and compact overdensity on top of the loose group.

The H I properties of HCG 44 have been studied by Williams & Rood (1987) using the Arecibo telescope, Williams, McMahon & van Gorkom (1991) and Verdes-Montenegro et al. (2001) with Very Large Array (VLA) data, van Driel et al. (2001) with the Nançay telescope, and Borthakur, Yun & Verdes-Montenegro (2010) using the Green Bank Telescope (GBT). These studies agree that galaxies in the group (all detected in H I except NGC 3193, which is claimed to be detected by van Driel et al. 2001 only) are gas poor relative to similar objects in the field. Verdes-Montenegro et al. (2001) estimate the detected H I mass to be $\sim 40$ per cent of the expected value for NGC 3187 and $\sim 10$ per cent for NGC 3185 and NGC 3190 (we revisit these estimates in Section 4.2). Borthakur et al. (2010) report the detection of some of the missing gas as they find excess H I emission in the GBT spectrum compared to the total VLA spectrum. Based on the strong similarity between the GBT and VLA H I profiles, they suggest that the excess gas is dynamically similar to the H I detected by the VLA within individual galaxies in the group.

We observed HCG 44 in H I as part of the multiwavelength ATLAS$^{3D}$ survey (Cappellari et al. 2011a). The H I observations of this survey were carried out with the Westerbork Synthesis Radio Telescope (WSRT) and presented in Serra et al. (2012). Those observations revealed a few small gas clouds around HCG 44. Serra et al. (2012) argue that the distribution of the clouds on the sky and in velocity is suggestive of a long, intragroup H I tail – the detected clouds being the densest clumps along the hypothetical tail. Here we present new, deeper WSRT observations performed as part of the ATLAS$^{3D}$ project to detect the H I tail itself.

## 3 OBSERVATIONS

### 3.1 H I interferometry

We observed HCG 44 for $6 \times 12$ h with the WSRT. We pointed the telescope at $\alpha_{2000} = 10^h 18^m 54.19$, $\delta_{2000} = 21^\circ 59' 30.5''$, which is a position between NGC 3193 and one of the H I clouds detected by Serra et al. (2012). We reduced the data in a standard way using the WSRT pipeline developed by Serra et al. (2012). The H I cube used for our analysis is made using robust = 0 weighting and 30 arcsec full width at half-maximum (FWHM) tapering. This results in a beam major and minor axis of 53.0 arcsec and 32.7 arcsec, respectively (PA = 6.5; the beam axes correspond to 6.4 and 4.0 kpc at the adopted distance, respectively). The cube has velocity resolution of 16 km s$^{-1}$ after Hanning smoothing. The noise is $\sigma = 0.22$ mJy beam$^{-1}$ in each channel, which corresponds to a formal 5$\sigma$ H I column density sensitivity of $1.1 \times 10^{19}$ cm$^{-2}$ per resolution element.

We use the source finder described in Serra et al. (2012) to detect emission in the H I cube. The finder looks for emission in the original cube and in cubes at different resolutions on the sky and/or in velocity. It flags all voxels above $+n\sigma$ and all voxels below $-n\sigma$ as

---

2 The GBT beam is $\sim 9$ arcmin.
emission and performs basic size filtering to reduce the noise in the
mask at each resolution. Here we use Gaussian filters of FWHM
25 and 50 arcsec on the sky and top-hat filters of width 16, 32,
64, 128, 256 and 384 km s\(^{-1}\) in velocity, and adopt \(n = 3\). We
use the resulting mask to build the total-H\(^i\) image (which we then
correct for the primary beam of the WSRT) and the H\(^i\) velocity
field (intensity-weighted mean) shown in Section 4. All H\(^i\) flux and
mass values reported in this paper are measured from the total-H\(^i\)
image and are therefore corrected for the primary beam.

3.2 Optical imaging

Deep optical imaging of HCG 44 was obtained with the MegaCam
camera mounted on the Canada–France–Hawaii Telescope (CFHT).
These observations were taken as part of the ATLAS\(^{10}\) project
(Cappellari et al. 2011a) with the goal of studying early-type galax-
ies’ morphological fine-structure as a relic of their formation (Duc
et al. 2011). We observed the HCG 44 field for 7 \(\times\) 345 s in both
g and \(r\) band applying offsets of \(\sim\)30 arcmin between consecutive
exposures. The seeing was 0.9 and 1.2 arcsec in \(g\) and \(r\) band,
respectively. We refer to Duc et al. (2011) for a full description
of the observations and data reduction. We note here that the ob-
serving strategy and data reduction procedures (ELIXIR–LSB software;
Cuillandre et al., in preparation) were chosen to detect stellar struc-
tures reaching \(\sim\)28.5 mag arcsec\(^{-2}\) in \(g\) band. However, the surface
brightness sensitivity may vary from field to field (and within a
given field), and we discuss the sensitivity of our image in more
detail in the following section.

4 RESULTS

4.1 H\(^i\) in the intragroup medium

Fig. 2 shows H\(^i\) constant-column-density contours overlaid on the
\(g\)-band CFHT/MegaCam image of HCG 44. The main result of our
new observation is the detection of a long H\(^i\) tail north of NGC
3193 (hereafter \(T_N\)), which is contained in the large, dashed red
circle in the figure. The tail consists of a low-column-density, \(\sim\)20-
arcmin-long component oriented north-west to south-east and a less
massive, southern extension. The latter is revealed by two small clouds
aligned north-east to south-west in the direction of NGC 3193. We also
confirm the Serra et al. (2012) detection of a smaller H\(^i\) complex
south-east of NGC 3193 (hereafter \(C_S\)), indicated by the
small, dashed red circle in the figure.

The tail \(T_N\) has no diffuse optical counterpart down to the surface
brightness sensitivity of the deep image shown in Fig. 2 (the same
result is obtained inspecting the \(r\)-band image). In Section 3.2,
we mentioned that the generic sensitivity of the CFHT/MegaCam
images taken as part of the ATLAS\(^{10}\) project is \(\sim\)28.5 mag arcsec\(^{-2}\)
in \(g\) band. In this particular case \(T_N\) lies on a relatively clean region
of the image, which shows no large-scale noise variations or bright
sources with the exception of two stars in the southern part of the tail.
The noise level in this region (excluding the two stars) corresponds
to a 1.5\(\sigma\) detection limit of 28.4 mag arcsec\(^{-2}\) within a circular
aperture of radius equal to the seeing, consistent with the generic
value given above. Such low detection level per resolution element is
sufficient considering that we are looking for emission over a
scale of many arcminutes.

Fig. 3 shows the velocity field of the detected H\(^i\). We find a
smooth variation of H\(^i\) recession velocity along \(T_N\). This is shown
more clearly by the position–velocity diagram in Fig. 4, which is
drawn along the path indicated by the black solid line in Fig. 3.
The clouds detected by Serra et al. (2012) and mentioned in Section
1 are therefore just the tip of the iceberg of a large, coherent, gas
distribution, as suggested in that paper. Gas in both \(T_N\) and \(C_S\) is
detected at velocities comparable to that of individual galaxies
in HCG 44, suggesting that the tails are parts of the group. \(T_N\)
has a total projected length of \(\sim\)220 kpc at the assumed distance
\(d_{\text{HCG 44}} = 25\) Mpc.

Table 2 lists the H\(^i\) flux of all detected objects shown in Fig. 2.
Table 2 also shows the fractional contribution \(f(\text{H}i)\) of each object
to the total-H\(^i\) flux of the group. Assuming that all galaxies in HCG
44 and the intragroup gas are roughly at the same distance, we
conclude that H\(^i\) in \(T_N\) and \(C_S\) amounts to \(\sim\)20 per cent of the total
neutral hydrogen mass of the compact group. This is as much H\(^i\)
as that found in NGC 3190 and almost half the H\(^i\) mass of NGC
3187. At the assumed distance \(d_{\text{HCG 44}}\) we estimate an H\(^i\) mass of
\(4.1 \times 10^8\) and \(1.2 \times 10^8\) \(M_\odot\) for \(T_N\) and \(C_S\), respectively.

We note that \(T_N\) stretches beyond the FWHM of WSRT’s primary
beam, which is 0.6 at the observing frequency of our data. Therefore,
it is possible that the gaseous tail is actually longer than shown
in Fig. 2 and we stop detecting H\(^i\) because of a decrease in the
telescope response. To check whether this is the case we make a
natural-weighting H\(^i\) cube from our WSRT data. This weighting
scheme produces a cube with very patchy noise and poorer image
quality, but the sensitivity is better than that of the main cube used
in this study. Fig. 5 shows a channel of the natural-weighting cube
at velocity 1248 km s\(^{-1}\) suggesting indeed that some additional H\(^i\)
emission may exist north-west of \(T_N\).

The existence of a north-western extension of \(T_N\) is confirmed by
data taken as part of the H\(^i\) Parkes All Sky Survey (HIPASS;
Barnes et al. 2001). The HIPASS cube\(^3\) of this sky area has noise of
\(\sim\)15 mJy beam\(^{-1}\), velocity resolution 18 km s\(^{-1}\) and beam FWHM
15.5 arcmin. The formal 5\(\sigma\) sensitivity per resolution element is
1.7 \(\times\) 10\(^{18}\) cm\(^{-2}\) for gas filling the beam. We run the source finder
described in Section 3.1 on the cube and show contours of the
detected H\(^i\) in Fig. 6 overlaid on an SDSS \(g\)-band image (we also
show the lowest WSRT H\(^i\) contour; see Fig. 2).

The figure shows two clear HIPASS detections at the location of
HCG 44 and NGC 3162 (a spiral galaxy \(\sim\)1.5 north-west of HCG
44). Both detections are listed in the northern HIPASS catalogue
by Wong et al. (2006) – objects HIPASS J1017+21 and HIPASS
J1013+22, respectively. The main part of \(T_N\) is visible only at a
tentative level in HIPASS (partly because the cube contains many
artefacts) and the HIPASS contours in Fig. 6 do not show it. On
the other hand, the HIPASS cube reveals additional emission (not
included in the catalogue of Wong et al. 2006) just north-west of \(T_N\),
which was missed by our WSRT data because of its large distance
from the pointing centre. The velocity of this emission is consistent
with the velocity field shown in Fig. 3. Including this emission, the
length of \(T_N\) becomes \(\sim\)300 kpc and its mass 5.2 \(\times\) 10\(^8\) \(M_\odot\). This
result is confirmed by the newly reduced HIPASS data, which have
significantly lower noise level and many less artefacts (Calabretta
et al., in preparation). In the new data the full tail is detected clearly
(HIPASS team, private communication).

Another interesting aspect revealed by the HIPASS data is that the
H\(^i\) tail stretches towards and possibly connects with NGC 3162
(this galaxy was already detected in H\(^i\) by e.g. van Driel et al. 2001).
NGC 3162, whose optical image we show in the inset of Fig. 6, has

\(^3\) Available at http://www.atnf.csiro.au/research/multibeam/release.
Figure 2. Constant-column-density HI contours overlaid on the g-band CFHT/MegaCam image. The colour bar on the right is in mag arcsec$^{-2}$. Contours represent $N$(HI) = $1.0 \times 10^{19} \times 3^n$ cm$^{-2}$ ($n = 0, 1, 2, 3$). Contours are coloured black to red, faint to bright. The black cross indicates the pointing centre of our WSRT observation and the dotted, black circle indicates the primary beam of the WSRT. Large and small dashed, red circles indicate the location of $T_N$ and $C_S$, respectively (see text). The beam of the HI image is shown on the bottom-left corner. Note that UGC 05574 (to the north-east) and SDSS J1017 (to the south-west) are not members of HCG 44.

HI recessional velocity of 1300 km s$^{-1}$, comparable to that of HCG 44 members, and is at a projected distance of $\sim$650 kpc from HCG 44 assuming the distance $d_{HCG44}$ along the line of sight. We discuss this finding in Section 5.

4.2 HI in galaxies: comparison to previous studies

All gas-rich galaxies in Fig. 2 were already known to host HI. Williams et al. (1991) observed this group with the VLA. They report HI fluxes systematically larger than those in Table 2. Namely, the flux of NGC 3185, NGC 3187, NGC 3190 and SDSS J1017 is a factor of 1.7, 1.3, 1.3 and 2.4 larger than our value, respectively.$^4$ For NGC 3185 and SDSS J1017 part of the difference may be due to the better column density sensitivity of the VLA data relative to the HI flux values are consistent with those reported by Williams & Rood (1987) using Arecibo data. They are adopted by Verdes-Montenegro et al. (2001) in their HI study of Hickson compact groups. They are 5–20 per cent larger than values in van Driel et al. (2001), which have however large uncertainty.

$^4$ These factors include a correction of the VLA fluxes for the primary beam of that telescope, which was not applied by Williams et al. (1991). The primary-beam correction is of $\sim$10 per cent for NGC 3185 and SDSS J1017, while it is negligible for the other two objects. Note that the VLA flux values are consistent with those reported by Williams & Rood (1987) using Arecibo data. They are adopted by Verdes-Montenegro et al. (2001) in their HI study of Hickson compact groups. They are 5–20 per cent larger than values in van Driel et al. (2001), which have however large uncertainty.
our WSRT data (by a factor of $\sim 1.7$; these galaxies are relatively far from the WSRT pointing centre as shown in Fig. 2). On the contrary, our data are slightly more sensitive at the location of NGC 3187 and NGC 3190 (10–20 per cent) so at least some of the flux difference must have a different cause.

Flux calibration is not the reason of this discrepancy as we have verified that our calibration is consistent with that of Williams et al. (1991) within $\pm 10$ per cent. Instead, we do find a relevant difference in the way we build the total-HI image. In our study we use the source finder of Serra et al. (2012), which includes in the HI image both positive and negative noise peaks (see Section 3.1). On the contrary, Williams et al. (1991) select only voxels above $+1.5\sigma$. This introduces a positive bias in the total-HI flux. Indeed, if we implement their detection criterion in our source finder we obtain fluxes a factor of 1.5, 1.1, 1.2 and 1.3 larger for the four galaxies mentioned above, respectively. This accounts at least partially for the difference in HI flux. Remaining differences may be explained by the better sensitivity of the VLA data at the location of NGC 3185 and SDSS J1017 (see above) and by the higher noise level of the VLA data at the location of NGC 3187 and NGC 3190, which would imply a higher level of the bias under discussion.

The HI morphology of galaxies in Fig. 2 is in agreement with the image presented by Williams et al. (1991). The main addition of our deeper image (besides the detection of $T_N$ and $C_S$) is a faint extension of the southern HI warp in NGC 3187. Furthermore, we set a column density limit of $\sim 2 \times 10^{19}$ cm$^{-2}$ on the HI bridge between NGC 3185 and NGC 3190, which was tentatively suggested by Williams et al. (1991) based on the VLA data (this limit takes into account that the two galaxies lay around the half-power point of the WSRT primary beam).
Table 2. H\textsc{i} detected with the WSRT in HCG 44.

<table>
<thead>
<tr>
<th>Object</th>
<th>(F(\text{H}\textsc{i})) (Jy km s(^{-1}))</th>
<th>(f(\text{H}\textsc{i}))</th>
<th>(M(\text{H}\textsc{i})) ((\times 10^8 \text{M}_\odot))</th>
</tr>
</thead>
<tbody>
<tr>
<td>(T_\text{S})</td>
<td>2.79</td>
<td>0.15</td>
<td>4.1</td>
</tr>
<tr>
<td>(C_\text{S})</td>
<td>0.82</td>
<td>0.05</td>
<td>1.2</td>
</tr>
<tr>
<td>NGC 3185</td>
<td>2.11</td>
<td>0.12</td>
<td>3.1</td>
</tr>
<tr>
<td>NGC 3187</td>
<td>8.22</td>
<td>0.46</td>
<td>12.0</td>
</tr>
<tr>
<td>NGC 3190</td>
<td>4.00</td>
<td>0.22</td>
<td>5.8</td>
</tr>
</tbody>
</table>

Galaxies not members of HCG 44

SDSS J1017 | 1.03 – |
UGC 05574 | 0.84 – |

Columns are listed as follows. (1) Object name. (2) Total \(\text{H}\textsc{i}\) flux. (3) Fraction of the total \(\text{H}\textsc{i}\) flux of HCG 44. (4) \(\text{H}\textsc{i}\) mass assuming a distance of \(d_{\text{HCG} 44} = 25\) Mpc for galaxies in HCG 44, and Hubble flow distance for galaxies outside HCG 44 (29 and 24 Mpc for SDSS J1017 and UGC 05574, respectively – see Table 1).

![Figure 4](https://example.com/figure4.png)

**Figure 4.** Position–velocity diagram of the \(\text{H}\textsc{i}\) emission along the path shown by the black solid line in Fig. 3. The origin of the horizontal axis is set to be at the northern end of the path. The bright emission at distance \(~45\) arcmin is \(\text{H}\textsc{i}\) in NGC 3187.

![Figure 5](https://example.com/figure5.png)

**Figure 5.** Contours of the \(\text{H}\textsc{i}\) emission in the natural-weighting cube at a recessional velocity of \(1248\) km s\(^{-1}\) overlaid on a grey-scale total-\(\text{H}\textsc{i}\) image of part of the tail \(T_\text{S}\). Grey contours are at \(~0.4\) mJy beam\(^{-1}\), black contours at \(+0.4\) and \(+0.8\) mJy beam\(^{-1}\). The figure shows possible additional emission north-west of \(T_\text{S}\). Note that this low-level emission was not cleaned (i.e. it was not deconvolved with the beam).

We note that the lower value of our \(M(\text{H}\textsc{i})\) estimates compared to values in Williams et al. (1991) would imply an even larger \(\text{H}\textsc{i}\) deficiency than that derived by Verdes-Montenegro et al. (2001) (see Section 2). We revisit the \(\text{H}\textsc{i}\) deficiency of galaxies in HCG 44 using the relation between \(M(\text{H}\textsc{i})\) and optical diameter \(D_{25}\), derived by Toribio et al. (2011). The isophotal major axis of NGC 3185, NGC 3187 and NGC 3190 is given in the SDSS Data Release 7 (DR7) catalogue and is 8.7, 6.5 and 13.4 kpc, respectively. Given these values, the \(\text{H}\textsc{i}\) mass predicted by the \(1/V_{\text{max}}\)-corrected \(M(\text{H}\textsc{i})\)–\(D_{25}\), relation in Toribio et al. (2011) is \(7.8, 5.4\) and \(13.5 \times 10^8 \text{M}_\odot\), respectively. Therefore, the detected mass of \(\text{H}\textsc{i}\) is just 4, 22 and 4 per cent of the expected value for the three galaxies (6\(\sigma, 3\sigma\) and \(6\sigma\) below the expected \(\text{H}\textsc{i}\) mass, respectively, where \(\sigma = 0.23\) dex is the rms residual of Toribio et al. 2011 relation).

5 DISCUSSION

The detection of a 300-kpc-long tail containing \(5 \times 10^8 \text{M}_\odot\) of \(\text{H}\textsc{i}\) in a group already observed by various authors with many different radio telescopes may seem surprising (see Section 2 for a summary of previous \(\text{H}\textsc{i}\) observations of HCG 44). Our result demonstrates that signatures of ongoing galaxy evolution inside groups can be truly elusive. It shows that although \(\text{H}\textsc{i}\) observations are unique in giving direct evidence of the fundamental role of group processes for galaxy evolution (as shown not only on individual systems but also on large, statistically representative samples; e.g. Serra et al. 2012), such observations need to be very sensitive and cover a large field if we want to gather a complete census of these events. This should be kept in mind when designing future \(\text{H}\textsc{i}\) surveys.

Besides the above general conclusion, our observations reveal for the first time direct evidence of gas stripping in HCG 44. This is interesting because galaxies in this group have long been known to be \(\text{H}\textsc{i}\) deficient and we may be unveiling the cause of at least part of the deficiency. It is therefore interesting to speculate on how the detected \(\text{H}\textsc{i}\) tail may have formed. How was the gas stripped, and from what galaxy? The goal of the present section is to explore possible answers to these two questions.

5.1 Formation of the \(\text{H}\textsc{i}\) tail: ram pressure or tidal stripping

The two possible mechanisms to form a gas tail within a group are ram pressure and tidal stripping. The former necessitates a dense medium and large relative velocity between the medium and the stripped galaxy. So far, X-ray observations have been unsuccessful in detecting the intragroup medium of HCG 44 (Mulchaey et al. 2003 using ROSAT data; Rasmussen et al. 2008 using Chandra and XMM–Newton data). Rasmussen et al. (2008) estimate that the density of the medium is \(n \sim 10^{-4} \text{cm}^{-3}\). Analytic calculations indicate that even at such low density some stripping might occur but they also show that this would be a fairly small effect (e.g. Rasmussen et al. 2008; Freeland, Sengupta & Croston 2010; Westmeier, Braun & Koribalski 2011). It therefore seems unlikely that ram pressure is responsible for creating such a prominent tail, longer than and as massive as tails detected in the centre of clusters (e.g. Oosterloo & van Gorkom 2005). Furthermore, no galaxy in the group shows the typical \(\text{H}\textsc{i}\) morphology caused by ram-pressure stripping – i.e. gas compressed against the stellar disc on one side of the galaxy and extending to larger radius on the opposite side (e.g. Kenney, van Gorkom & Vollmer 2004; Chung et al. 2007, 2009). We conclude that ram pressure is an unlikely explanation for the formation of \(T_\text{S}\).

The alternative hypothesis is that \(T_\text{S}\) was created by tidal interaction within the group. Tidal forces act on both stars and gas and it...
may therefore seem surprising that no diffuse stellar light is detected at the location of $T_N$ even in the very deep CFHT/MegaCam image shown in Fig. 2. In fact, observations (e.g. Davies et al. 2004) and simulations (e.g. Bekki et al. 2005; Duc & Bournaud 2008; Michel-Dansac et al. 2010) both show that starless tidal tails can develop following gravitational interaction and depending on the relative distribution of gas and stars in the stripped galaxy. Therefore, tidal interaction is a viable mechanism to create the detected tail and we explore this possibility in the rest of this section.

5.2 What stripped galaxy?

The first question we attempt to answer is what galaxy the $H_1$ was stripped from. The results described in Section 4 suggest two alternative hypotheses. One possibility is that the tail was stripped from NGC 3162 as it flew by the group at high speed (see Fig. 6). This galaxy is currently at a projected distance of $\sim 650$ kpc from HCG 44, which it could have covered in $\sim 3$ Gyr if it went through the group at a velocity of 200 km s$^{-1}$ on the plane of the sky. The optical morphology of this galaxy may support this hypothesis as the stellar disc is lopsided, indicating that it might have been perturbed recently (see inset in Fig. 6). However, lopsidedness is a relatively common phenomenon in spiral galaxies and is no definite proof of an interaction between NGC 3162 and HCG 44. The newly reduced HIPASS data (Calabretta et al., in preparation) hint to a possible H$_1$ bridge between $T_N$ and NGC 3162 which, if confirmed, would be a strong clue in favour of this hypothesis. This will be explored further in a future paper presenting the new data.

The other possibility is that the observed $H_1$ tail originated from one of the galaxies currently within HCG 44, and NGC 3162 played no significant role. Optical imaging shows that some of the members of HCG 44 have experienced recent tidal interaction (see Fig. 1), and $T_N$ may have formed as part of the same process. The most obvious candidate for tidal stripping within the group might be NGC 3187. This galaxy exhibits a strong tidal distortion visible as an S-shaped warp and $T_N$ could be seen as an extension of the north-east side of the warp (see Fig. 2; note that any connection between $T_N$ and NGC 3187 would have to be at column density below $\sim 10^{19}$ cm$^{-2}$). This is confirmed by the fact that the $H_1$ velocity on the south-west end of $T_N$ is similar to that on the...
north-east side of NGC 3187 (Figs 3 and 4). Another interesting point is that gas in T_N amounts to almost half of the total-H\(_{I}\) mass of NGC 3187. Therefore, tidal stripping would provide a possible explanation for the H\(_{I}\) deficiency of this galaxy (see Section 4.2). We speculate on the details of the stripping process below.

### 5.3 Tidal stripping of NGC 3187

Numerous previous authors argued that NGC 3187 might be interacting with NGC 3190. While this is possible, it is unlikely that such interaction is responsible for the formation of T_N. The reason is that it would be difficult to explain the gap between H\(_{I}\) in NGC 3187 and T_N. Furthermore, compared to previously known cases of very long tails induced by \(~\)major galaxy interaction and merging (e.g. Mirabel, Lutz & Maza 1991; Duc et al. 2011), the length of the tail – at least 300 kpc – seems too extreme to be caused by this particular galaxy pair.

A mechanism to form very long, low-column-density H\(_{I}\) tails (or rings) like that in HCG 44 was proposed by Bekki et al. (2005). The main ingredient of their model is the tidal interaction of a low-surface-brightness disc (in this case NGC 3187) with the gravitational potential of a group of galaxies. This interaction would strip the outer part of the disc (which are often observed to be essentially starless) and distribute the stripped material on a long tail or ring.

We perform a basic test of this mechanism and estimate the timescale of the stripping process by building a 3D toy model for the orbit of NGC 3187 around HCG 44. The model is constrained by the observed trajectory of NGC 3187 (traced by the H\(_{I}\) tail and the current position of the galaxy) and its velocity relative to the group. In this model we assume the group potential to be fixed and Keplerian. We take the centre of the potential to be at the position of NGC 3193 and place it at coordinates \([x, y, z] = [0, 0, 0]\) (we discuss the orientation of the sky relative to these three axes below). We further assume that NGC 3187 follows a parabolic orbit parallel to the \(xy\) plane (and therefore at constant \(z = z_0\)).

For \(z_0 = 0\) the orbit is determined only by the total mass of the system \(M_{tot}\) and the focal length of the parabola \(f\). However, in this model we assume that the orbital plane is displaced relative to the centre of the group (\(z_0 \neq 0\); we show below that this is necessary for the parabola to fit the data). Under this assumption NGC 3187 should move also along the \(z\) axis but we neglect this effect. Such additional motion would be little constrained by the available data and would not in any case have a significant impact on the timescale of the orbit, which is our primary goal. Therefore, the orbit remains perfectly planar (this approximation becomes increasingly wrong as the \(xy\) distance of NGC 3187 from the group centre gets closer to the value of \(z_0\)).

The toy model described above has only three free parameters: \(M_{tot}\), \(f\) and \(z_0\). We fix \(M_{tot} = 5 \times 10^{12} M_{\odot}\). We vary \(f\) between 50 and 200 kpc, and \(z_0\) between 0 and 200 kpc in an attempt to reproduce the observations. Furthermore, in order to compare the model to the data we need to project it on the sky. This introduces additional degrees of freedom. We make use of 3D visualization to find a favourable projection.

Fig. 7 shows a possible orbit of NGC 3187, which reproduces the location of the H\(_{I}\) tail and that of the galaxy at once. The orbit is obtained with \(f = 75\) kpc and \(z_0 = 100\) kpc, and is viewed at an inclination of 60° away from face-on (note that this is not the result of a formal fit and is therefore just one possible solution of the problem). An important feature of the model is that the predicted velocity of NGC 3187 relative to the group centre (i.e. NGC 3193) is 235 km s\(^{-1}\) along the line of sight. This is a good match to the observed value of 210 km s\(^{-1}\) (see Table 1). Therefore, this simple model can explain the anomalous, high velocity of NGC 3187 relative to other members of HCG 44.

Within the context of this simple model NGC 3187 passed the vertex of the orbit \(~\)400 Myr ago and was at the current location of the north-west end of the H\(_{I}\) tail (revealed by the HIPASS data shown in Fig. 6) \(~\)900 Myr ago. This suggests that the galaxy may have been stripped of the H\(_{I}\) in \(T_N\) (\(~\)1/3 of its initial H\(_{I}\) mass) within a time interval shorter than 1 Gyr. It is interesting that this time-scale is consistent with the \(~\)2 Gyr time-scale for quenching of star formation in galaxy groups recently derived by Rasmussen et al. (2012b).

There are some important differences between HCG 44 and the systems simulated by Bekki et al. (2005). First, their simulations predict the existence of a leading tail which, in the case of NGC 3187, should start at the southern tip of the warp, bend towards north-west and eventually join \(T_N\) to form an intragroup H\(_{I}\) ring. This is not observed. Bekki et al. model is, however, an idealized one. In reality this process occurs in the presence of many other galaxies. It is possible that such second tail was destroyed by an interaction with NGC 3190 (which is tidally disturbed) and gas was instead dispersed towards the location of \(C_N\). This idea is supported by both the H\(_{I}\) and optical morphology of NGC 3187’s southern tidal tail. Fig. 2 shows that this tail broadens and bends towards NGC 3190, giving support to the idea that the two galaxies are interacting. Therefore, it is possible that also gas in \(C_N\) once belonged to NGC 3187, and the above estimate on the amount of gas stripped from this galaxy should be regarded as a lower limit.

Another difference is that in Bekki et al. model the stripped galaxy is not tidally distorted, unlike NGC 3187. Again, interaction with NGC 3190 may explain the difference (note that, as we have argued above, this interaction is unlikely to be the formation mechanism of \(T_N\) itself).

We stress that we view this model as a very simple (but nevertheless useful) description of a possible formation mechanism for the detected tail. However, the system is complex: the group
membership is still to be established and a number of galaxies might be playing a relevant role, including the very distant NGC 3162 (see Fig. 6 and Section 5.2). As a consequence, substantial differences between the data and idealized simulations like that of Bekki et al. (2005) or simple models like the parabolic trajectory presented above are not a surprise. More detailed modelling is beyond the scope of this paper (for example, a realistic mass model of the group would be needed, and the line-of-sight velocity of gas in $T_g$ should be used as a constraint for the model – taking into account also rotation within the stripped galaxy). Such modelling or the exploration of different formation mechanisms for the tail (e.g. interaction between NGC 3162 and a member of HCG 44 within the group potential, as in the models by Bekki, Koribalski & Kilborn 2005) would strongly benefit from having deep H$\alpha$ observations over a larger field.

6 CONCLUSIONS

We have presented deep H$\alpha$ and optical imaging of the galaxy group HCG 44 obtained with the WSRT and CFHT/MegaCam, respectively, as part of the ATLAS$^{3D}$ project. We detect a long intragroup H$\alpha$ tail and additional intragroup gas which, together, amount to 20 per cent of the H$\alpha$ mass of the group. Combining these data with archive HIPASS observations we find that the main tail contains $\sim 5 \times 10^8 M_\odot$ of H$\alpha$ and is $\sim 300$ kpc long assuming a distance of 25 Mpc. The tail has no diffuse optical counterpart down to $\sim 28.5$ mag arcsec$^{-2}$ in $g$ band.

We discuss viable formation mechanisms for the H$\alpha$ tail. Based on the available data (including X-ray imaging), it is unlikely that the tail is caused by ram-pressure stripping. Instead, it is possible that the H$\alpha$ was stripped from NGC 3187, a member of HCG 44, by the group tidal field. We present a simple model for the orbit of the stripped galaxy through the group. Within this model, NGC 3187 has been stripped of $\frac{1}{3}$ of its initial gas mass in less than 1 Gyr. This is consistent with recent estimates of the time-scale for quenching of star formation in galaxy groups ($\sim 2$ Gyr).

The proposed model is intentionally simple and it is possible that other processes have contributed to shaping the properties of galaxies in the group (e.g. tidal interaction between group members and, possibly, some ram-pressure stripping). Another possibility suggested by HIPASS data is that H$\alpha$ in the tail was stripped from NGC 3162, a spiral galaxy now 650 kpc from the group. Future work will investigate this possibility using H$\alpha$ data of better quality than those available at the moment.

Regardless of the precise formation mechanism, the detected H$\alpha$ tail is the first, direct evidence of gas stripping in HCG 44, a group long known to be deficient of H$\alpha$. Our result highlights the importance of group processing as a driver of galaxy evolution, but also the observational challenge that has to be overcome in order to detect signatures of these processes. Sensitive H$\alpha$ observations over a large field are needed to gather a complete census of this kind of events in the local Universe.

ACKNOWLEDGMENTS

PS acknowledges useful discussions with T. van Albada, J. van Gorkom, R. Sancisi, S. Tomnesen and C. Toribio.

MC acknowledges support from a Royal Society University Research Fellowship. This work was supported by the rolling grants ‘Astrophysics at Oxford’ PP/E001114/1 and ST/H002456/1 and visitors grants PPA/V/S/2002/00553, PP/E001564/1 and ST/H504862/1 from the UK Research Councils. RLD acknowledges travel and computer grants from Christ Church, Oxford and support from the Royal Society in the form of a Wolfson Merit Award 502011.K502/jd. RLD also acknowledges the support of the ESO Visitor Programme which funded a 3-month stay in 2010. SK acknowledges support from the Royal Society Joint Projects Grant JP0869822. RMM is supported by the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., on behalf of the international Gemini partnership of Argentina, Australia, Brazil, Canada, Chile, the United Kingdom and the United States of America. LM-D acknowledges support from the Lyon Institute of Origins under grant ANR-10-LABX-66. TN and MB acknowledge support from the DFG Cluster of Excellence ‘Origin and Structure of the Universe’. MS acknowledges support from an STFC Advanced Fellowship ST/F009186/1.

PS is a NWO/Veni fellow. (TAD) The research leading to these results has received funding from the European Community’s Seventh Framework Programme (FP7/2007-2013/) under grant agreement No 229517. MB has received, during this research, funding from the European Research Council under the Advanced Grant Promotion Number 267399-Momentum. The authors acknowledge financial support from ESO. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This paper is based on observations obtained with the Westerbork Synthesis Radio Telescope, which is operated by the ASTRON (Netherlands Foundation for Research in Astronomy) with support from the Netherlands Foundation for Scientific Research NWO, and with MegaPrime/MegaCam, a joint project of CFHT and CEA/DAPNIA, at the CFHT, which is operated by the National Centre for Scientific Research of France and the University of Hawaii.

REFERENCES

Garcia A. M., 1993, A&AS, 100, 47

379 A giant H$\alpha$ tail in HCG 44